## SCOPE ON THE SKIES

## Tracking planets around the Sun

by Bob Riddle

Since this series about the solar system began last September, the six visible planets, as well as the other two planets too distant to be seen with unaided eyes, have traveled along their respective orbital paths. In the process, the planets have, in some instances, shifted from visibility during the evening in the west, to being visible in the morning skies in the east. This is occurring regularly and the changing locations of the planets can be tracked in a variety of ways using any of several celestial coordinate systems.

In earlier columns, the celestial coordinate system of hour circles of right ascension and degrees of declination was introduced along with the use of an equatorial star chart (see SFA Star Charts in Resources). This system shows the planets' motion relative to the ecliptic, the apparent path the Sun follows during the year. An alternate system, using heliocentric longitude, places the viewer in an above-the-solar-system viewpoint, tracking the planets in their near-circular orbits around the Sun.

## Heliocentric longitude

This is a coordinate system where planets are plotted on polar graph paper (see Resources) using degrees ranging from 0 to 360 . The Sun is located at the $0^{\circ}$ point and does not move, while the planets move in increasing longitude counterclockwise from the $0^{\circ}$ point. Using the data in

Figure 2, students can plot the positions for the planets over a three-month period, approximately the orbital period or length of one Mercurian year. For the outermost planets, the longitudes are given using more precise values, as these planets move less than $1^{\circ}$ each month. One degree in the sky is approximately twice the angular diameter of the full Moon (29-33 minutes of arc).

As students plot the planet positions, they should keep in mind that the orbits they are plotting are on circles, while all of the planets actually have orbits that are somewhat elliptical. Students can use Figure 1 to compare the shapes of the orbits. Eccentricity is a measure of how circular or elliptical the shape of an orbit is using a scale from 0 to 1 where 0 is a circle and 1 is a straight line. This is an appropriate time to review or introduce Kepler's laws of planetary motion.

To continue the plotting of planetary heliocentric longitudes, download or print the monthly coordinates from the Planet Watch website (see Resources). An alternate source that gives more flexibility in its application is the freeware program Interactive Computer Ephemeris (ICE). It can be downloaded from the link provided (see Resources) and used to generate heliocentric longitudes, as well as other astronomical data. The program runs in a DOS window on a PC and is a bit confusing at first use; once it is

## FIGURE 1 Planet orbital data

|  | Mercury | Venus | Earth | Mars | Jupiter | Saturn | Uranus | Neptune |
| :--- | :---: | :--- | :--- | :--- | :--- | :--- | :--- | :--- | :--- |
| Distance from Sun <br> (million miles) | 36.0 | 67.2 | 93.0 | 141.6 | 483.8 | 890.8 | $1,784.8$ | $2,793.1$ |
| Perihelion <br> (million miles) | 28.6 | 66.8 | 91.4 | 128.4 | 460.1 | 840.4 | $1,703.4$ | $2,761.6$ |
| Aphelion <br> (million miles) | 43.4 | 67.7 | 94.5 | 154.9 | 507.4 | 941.1 | $1,866.4$ | $2,824.5$ |
| Orbital period <br> (Earth days) | 88.0 | 224.7 | 365.2 | 687.0 | 4,331 | 10,747 | 30,589 | 59,800 |
| Orbital eccentricity | 0.205 | 0.007 | 0.017 | 0.094 | 0.049 | 0.057 | 0.046 | 0.011 |


| FIGURE 2 | Heliocentric coordinates and avera |  |
| :---: | :---: | :---: |
| Mercury |  |  |
| Date | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| March 1 | 227 | 0.450 |
| March 31 | 314 | 0.416 |
| April 10 | 354 | 0.361 |
| April 30 | 110 | 0.316 |
| May 10 | 163 | 0.364 |
| May 30 | 233 | 0.456 |
| Venus |  |  |
| Date | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| March 1 | 279 | 0.727 |
| March 31 | 326 | 0.7280 |
| April 10 | 342 | 0.7275 |
| April 30 | 14 | 0.7255 |
| May 10 | 30 | 0.7242 |
| May 30 | 62 | 0.7215 |
| Earth |  |  |
| Date | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| March 1 | 160 | 0.9908 |
| March 31 | 190 | 0.9990 |
| April 10 | 200 | 1.001 |
| April 30 | 220 | 1.007 |
| May 10 | 229 | 1.009 |
| May 30 | 248 | 1.013 |
| Mars |  |  |
| Date | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| March 1 | 124 | 1.639 |
| March 31 | 137 | 1.656 |
| April 10 | 141 | 1.660 |
| April 30 | 150 | 1.665 |
| May | 154 | 1.665 |
| May 30 | 163 | 1.664 |

set up, it will calculate useful data about planets for students to plot and study.

## March equinox

On March 20, 0148 EST (1:48 a.m.), the Sun will reach the astronomical coordinates of 0 hours and 0 degrees. This is one of two days during the year where, in general, the Sun rises due east and sets due west, giving rise to an

| Jupiter |  |  |
| :--- | :--- | :--- |
| Date | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| March 1 | 276 | 5.225 |
| March 31 | 279 | 5.214 |
| April 30 | 281 | 5.203 |
| May 30 | 284 | 5.192 |
|  |  |  |
| Saturn |  |  |
| Date | Longitude ( $\left.{ }^{\circ}\right)$ | Distance (AU) |
| March 1 | 155 | 9.281 |
| March 31 | 156 | 9.290 |
| April 30 | 157 | 9.298 |
| May 30 | 158 | 9.306 |
|  |  |  |
| Uranus | Longitude ( ${ }^{\circ}$ ) | Distance (AU) |
| Date | 348 | 20.096 |
| March 1 | 348 | 20.096 |
| March 31 | $34912^{\prime}$ | 20.097 |
| April 30 | $34931^{\prime}$ | 20.097 |
| May 30 |  |  |
|  |  |  |
| Neptune | Longitude ( $\left.{ }^{\circ}\right)$ | Distance (AU) |
| Date | $32150^{\prime}$ | 30.041 |
| March 1 | $32200^{\prime}$ | 30.040 |
| March 31 | $32211^{\prime}$ | 30.039 |
| April 30 | $32222^{\prime}$ | 30.039 |
| May 30 |  |  |

equal amount of daylight and night, or an equinox. This is also the day marking a change of seasons for both the Northern and Southern Hemispheres. Thinking globally, an equinox should be known by the month during which it occurs rather than the traditional and Northern Hemisphere term. Instead of the vernal or spring equinox, we should refer to it as the March equinox.

## Visible planets

Mercury will be visible low over the eastern horizon before sunrise. Use the much brighter-appearing Venus to help locate Mercury.
Venus will be visible low over the eastern horizon before the sunrise. Watch for the much dimmer-appearing Mercury to lie to the right of Venus early in the month

## FIGURE 3

Heliocentric positions of planets for March 1, 2008

and to the left of Venus by the end of the month.
Mars will be over the southern horizon at sunset and will be visible the rest of the night, setting a few hours after midnight.
Jupiter will rise about two hours before sunrise and will be easily seen shining brightly over the southeastern horizon.
Saturn will rise after sunset and will be visible the remainder of the night. By sunrise, Saturn will be low over the western horizon.

## Celestial events

3/2 Moon west of Jupiter
$3 / 3 \quad$ Moon east of Jupiter
Mercury at west elongation
3/5 Thin crescent Moon near Mercury and Venus
3/7 New Moon
3/9 Spring forward-Begin Daylight Saving Time Uranus in conjunction with Sun
3/12 Cassini flyby of Enceladus Mars passes M35
3/14 First quarter Moon Moon near Mars
3/18 Moon near Saturn and Regulus
3/20 March equinox (0548 UT)
3/21 Full Moon

3/25 Cassini flyby of Titan
$3 / 29$ Last quarter Moon
3/30 Moon near Jupiter

## Questions for students

1. How many degrees does the Earth move along its orbital path each day? Explain how this is different from the amount any of the outer planets move each day. (The Earth moves nearly $1^{\circ}\left[0.9856^{\circ}\right]$ each day in its orbit around the Sun. Because an outer planet, one further from the Sun, takes longer to orbit the Sun than one closer to the Sun, the Earth moves more quickly, and thus travels more along its orbit than a more distant outer planet.)
2. Using Figure 1, how would you know which of the eight planets has the greatest range in distance from the Sun? (The eccentricity of an orbiting object is a measure of how circular the orbit is. The values for eccentricity range from circular, 0 , to a straight line, 1. Mercury has the greatest eccentricity of the planets with a value of 0.205 .)
3. Use the internet to research Johannes Kepler and the laws of planetary motion he developed. Which of the three Kepler's laws applies to the varying speed that planets have as their distance from the Sun increases? (Kepler's second law, sometimes referred to as the law of equal areas, is a description of the speed each planet will move as it orbits the Sun. Planets move more quickly when close to the Sun and more slowly when further away.)

## Resources

Daylight Saving Time-http://geography.about.com/cs/ daylightsavings/a/dst.htm
Heliocentric calculator-http://nssdc.gsfc.nasa.gov/space/ helios/planet.html
ICE software-www.astro.uio.no/ita/TNP/ice/ice.html
Planet watch-http://currentsky.com
Polar graph-paper maker-www.incompetech.com/graph paper/polar
Riddle, B. 2006. Scope on the skies: The equinox. Science Scope 29 (6): 78-79.
SFA star charts-www.midnightkite.com/starcharts.html
Sun Shadow Investigation Project-http://sunship. currentsky.com
The Eggquinox-http://currentsky.com/articles/eggquinox

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